

Neutrino properties from high energy astrophysical neutrinos

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It is shown how high energy neutrino beams from very distant sources can be utilized to learn about some properties of neutrinos such as lifetimes, mass hierarchy, etc. Furthermore, even mixing elements such as U_{e3} and the CPV phase in the neutrino mixing matrix can be measured in principle. Pseudo-Dirac mass differences as small as $10^{-18}eV^2$ can be probed as well.

1. Introduction

We make two basic assumptions which are reasonable. The first one is that distant neutrino sources (e.g. AGN's and GRB's) exist; and furthermore with detectable fluxes at high energies (upto and beyond PeV). The second one is that in the not too far future, very large volume, well instrumented detectors of sizes of order of KM3 and beyond will exist and be operating; and furthermore will have (a) reasonably good energy resolution and (b) good angular resolution ($\sim 1^\circ$ for muons).

2. Neutrinos from Astrophysical Sources

If these two assumptions are valid, then there are a number of uses these detectors can be put to[1]. In this talk I want to focus on those that enable us to determine some properties of neutrinos: namely, probe neutrino lifetimes to $10^4s/eV$ (an improvement of 10^8 over current bounds), pseudo-Dirac mass splittings to a level of $10^{-18}eV^2$ (an improvement of a factor of 10^6 over current bounds) and potentially even measure quantities such as U_{e3} and the phase δ in the MNSP matrix[2].

3. Astrophysical neutrino flavor content

In the absence of neutrino oscillations we expect a very small ν_τ component in neutrinos from astrophysical sources. From the most discussed and the most likely astrophysical high energy neutrino sources[3] we expect nearly equal numbers of particles and anti-particles, half as many ν'_e 's as ν'_μ 's and virtually no ν'_τ 's. This comes about simply because the neutrinos are thought to originate in decays of pions (and kaons) and subsequent decays of muons. Most astrophysical targets are fairly tenuous even compared to the Earth's atmosphere, and would allow for full muon decay in flight. There are some predictions for flavor independent fluxes from cosmic defects and exotic objects such as evaporating black holes. Observation of tau neutrinos from these would have great importance. A conservative estimate[4] shows that the prompt ν_τ flux is very small and the emitted flux is close to the ratio $1 : 2 : 0$. The flux ratio of $\nu_e : \nu_\mu : \nu_\tau = 1 : 2 : 0$ is certainly valid for those AGN models in which the neutrinos are produced in beam dumps of photons or protons on matter, in which mostly pion and kaon decay (followed by the decay of muons) supply the bulk of the neutrino flux.

Depending on the amount of prompt ν -flux

due to the production and decay of heavy flavors, there could be a small non-zero ν_τ component present. There are also possible scenarios in which the muons lose energy in matter or in strong magnetic fields[5], in which case the initial flux mixture becomes $\nu_e : \nu_\mu : \nu_\tau = 0 : 1 : 0$.

4. Effect of Oscillations

The current knowledge of neutrino masses and mixings can be summarized as follows[6]. The mixing matrix is given to a good approximation by

$$U = \begin{pmatrix} c & s & \epsilon \\ s/\sqrt{2} & c/\sqrt{2} & 1/\sqrt{2} \\ s/\sqrt{2} & c/\sqrt{2} & 1/\sqrt{2} \end{pmatrix} \quad (1)$$

where $c = \cos \theta$, $s = \sin \theta$ with θ the solar mixing angle given by about 32° and $\epsilon = U_{e3} < 0.17$ limited by the CHOOZ bound. The mass spectrum has two possibilities; normal or inverted, and with the mass differences given by $\delta m_{32}^2 \sim 2.10^{-3} eV^2$ and $\delta m_{21}^2 \sim 7.10^{-5} eV^2$. Since $\delta m^2 L/4E$ for the distances to GRB's and AGN's (even for energies upto and beyond PeV) is very large ($> 10^7$) the oscillations have always averaged out and the conversion and survival probabilities are given by

$$P_{\alpha\beta} = \sum_i |U_{\alpha i}|^2 |U_{\beta i}|^2 \quad (2)$$

$$P_{\alpha\alpha} = \sum_i |U_{\alpha i}|^4 \quad (3)$$

Assuming no significant matter effects enroute, the mixing matrix in Eq. (1) leads to a propagation matrix P, given by:

$$P = \begin{pmatrix} 1 - S/2 & S/4 & S/4 \\ S/4 & 1/2 - S/8 & 1/2 - S/8 \\ S/4 & 1/2 - S/8 & 1/2 - S/8 \end{pmatrix} \quad (4)$$

where S stands for $\sin^2(2\theta)$. As is obvious, for any value of the solar mixing angle, P converts a flux ratio of $\nu_e : \nu_\mu : \nu_\tau = 1 : 2 : 0$ into one of

$1 : 1 : 1$. Hence the flavor mix expected at arrival is simply an equal mixture of ν_e, ν_μ and ν_τ as was observed long ago[4,7]. An initial flavor mix of $\nu_e : \nu_\mu : \nu_\tau = 0 : 1 : 0$ is converted by oscillations into one of about $1/2 : 1 : 1$. There are several other ways, arising from intrinsic properties of neutrinos, by which the flavor mix can change from the canonical $1 : 1 : 1$ figure. One in particular, which gives rise to striking signatures, is the decay of neutrinos[8]. Before discussing the other possibilities, let me consider the case of neutrino decay.

5. Neutrino Decay[9]

We now know that neutrinos have non-zero masses and non-trivial mixings, based on the evidence for neutrino mixings and oscillations from the data on atmospheric, solar and reactor neutrinos.

If this is true, then in general, the heavier neutrinos are expected to decay into the lighter ones via flavor changing processes. The only questions are (a) whether the lifetimes are short enough to be phenomenologically interesting (or are they too long?) and (b) what are the dominant decay modes.

Throughout the following, to be specific, I will assume that the neutrino masses are at most of order of eV. Since we are interested in decay modes which are likely to have rates (or lead to lifetimes) which are phenomenologically interesting, we can rule out several classes of decay nodes.

First, consider radiative decays, such as $\nu_i \rightarrow \nu_j + \gamma$. Since the experimental bounds on μ_{ν_i} , the magnetic moments of neutrinos, come from reactions such as $\nu_e e \rightarrow e \nu$ which are not sensitive to the final state neutrinos; the bounds apply to both diagonal as well as transition magnetic mo-

ments and so can be used to limit the corresponding lifetimes. The bounds should really be on mass eigenstates[10], but since the mixing angles are large, it does not matter much. The current bounds are[11]:

$$\begin{aligned}\tau_{\nu_e} &> 5.10^{18} \text{ sec} \\ \tau_{\nu_\mu} &> 5.10^{16} \text{ sec} \\ \tau_{\nu_\tau} &> 2.10^{11} \text{ sec}\end{aligned}\tag{5}$$

In the above limits the first one gives a bound for the τ_{ν_1} , whereas the second one gives the bound for both τ_{ν_2} as well as τ_{ν_3} since the mixing is essentially maximal.

There is one caveat in deducing these bounds. Namely, the form factors are evaluated at $q^2 \sim O(eV^2)$ in the decay matrix elements whereas in the scattering from which the bounds are derived, they are evaluated at $q^2 \sim O(MeV^2)$. Thus, some extrapolation is necessary. It can be argued that, barring some bizarre behaviour, this is justified[12].

An invisible decay mode with no new particles is the three body decay $\nu_i \rightarrow \nu_j \nu_j \bar{\nu}_j$. Even at the full strength of Z coupling, this yields a lifetime of $2.10^{34}s$, far too long to be of interest. There is an indirect bound from Z decays which is weaker but still yields $2.10^{30}s$ [13].

Thus, the only decay modes which can have interestingly fast decays rates are two body modes such as $\nu_i \rightarrow \nu_j + x$ and $\nu_i \rightarrow \bar{\nu}_j + x$ where x is a very light or massless particle, e.g. a Majoron.

The only possibility for fast invisible decays of neutrinos seems to lie with Majoron or Majoron-like models[9]. There are two classes of models; the I=1 Gelmini-Roncadelli[14] majoron and the I=0 Chikasige-Mohapatra-Peccei[15] majoron. In general, one can choose the majoron to be a mixture of the two; furthermore the coupling can be

to flavor as well as sterile neutrinos. The effective interaction is of the form:

$$\bar{\nu}_\beta^c (a + b\gamma_5) \nu_\alpha J \tag{6}$$

giving rise to decay:

$$\nu_\alpha \rightarrow \bar{\nu}_\beta \text{ (or } \nu_\beta) + J \tag{7}$$

where J is a massless $J = 0, L = 2$ particle; ν_α and ν_β are mass eigenstates which may be mixtures of flavor and sterile neutrinos. Models of this kind which can give rise to fast neutrino decays have been discussed[16]. These models are unconstrained by μ and τ decays which do not arise due to the $\Delta L = 2$ nature of the coupling. The I=1 coupling is constrained by the bound on the invisible Z width; and requires that the Majoron be a mixture of I=1 and I=0[17]. The couplings of ν_μ and ν_e (g_μ and g_e) are constrained by the limits on multi-body π , K decays $\pi \rightarrow \mu\nu\nu\nu$ and $K \rightarrow \mu\nu\nu\nu$ and on $\mu - e$ universality violation in π and K decays[18], but not sufficiently strongly to rule out fast decays.

There are very interesting cosmological implications of such couplings. The details depend on the spectrum of neutrinos and the scalars in the model. For example, if all the neutrinos are heavier than the scalar; the relic neutrino density vanishes today, and the neutrino mass bounds from CMB and large scale structure are no longer operative, whereas future measurements in the laboratory might find a non-zero result for a neutrino mass [19]. If the scalars are heavier than the neutrinos, there are signatures such as shifts of the n th multipole peak (for large n) in the CMB [20]. There are other implications as well, such as the number of relativistic degrees of freedom (or effective number of neutrinos) being different at the BBN and the CMB eras. The additional degrees

of freedom should be detectable in future CMB measurements.

Direct limits on such decay modes are also very weak. Current bounds on such decay modes are as follows. For the mass eigenstate ν_1 , the limit is about

$$\tau_1 \geq 10^5 \text{ sec}/eV \quad (8)$$

based on observation of $\bar{\nu}_e s$ from SN1987A [21] (assuming CPT invariance). For ν_2 , strong limits can be deduced from the non-observation of solar anti-neutrinos in KamLAND[22] but only in the case when the coupling is to ν_1 . In the most general case, an analysis of solar neutrino data[23] leads to a bound given by:

$$\tau_2 \geq 10^{-4} \text{ sec}/eV \quad (9)$$

For ν_3 , in case of normal hierarchy, one can derive a bound from the atmospheric neutrino observations of upcoming neutrinos[24]:

$$\tau_3 \geq 10^{-10} \text{ sec}/eV \quad (10)$$

The strongest lifetime limit is thus too weak to eliminate the possibility of astrophysical neutrino decay by a factor about $10^7 \times (L/100 \text{ Mpc}) \times (10 \text{ TeV}/E)$. Some aspects of the decay of high-energy astrophysical neutrinos have been considered in the past. It has been noted that the disappearance of all states except ν_1 would prepare a beam that could in principle be used to measure elements of the neutrino mixing matrix, namely the ratios $U_{e1}^2 : U_{\mu 1}^2 : U_{\tau 1}^2$ [25]. The possibility of measuring neutrino lifetimes over long baselines was mentioned in Ref.[26], and some predictions for decay in four-neutrino models were given in Ref.[27]. We have shown that the particular values and small uncertainties on the neutrino mixing parameters allow for the first time very distinctive signatures of the effects of neutrino decay

on the detected flavor ratios. The expected increase in neutrino lifetime sensitivity (and corresponding anomalous neutrino couplings) by several orders of magnitude makes for a very interesting test of physics beyond the Standard Model; a discovery would mean physics much more exotic than neutrino mass and mixing alone. As shown below, neutrino decay because of its unique signature cannot be mimicked by either different neutrino flavor ratios at the source or other non-standard neutrino interactions.

A characteristic feature of decay is its strong energy dependence: $\exp(-Lm/E\tau)$, where τ is the rest-frame lifetime. For simplicity, we will assume that decays are always complete, i.e., that these exponential factors vanish. The assumption of complete decay means we do not have to consider the distance and intensity distributions of sources. We assume an isotropic diffuse flux of high-energy astrophysical neutrinos, and can thus neglect the angular deflection of daughter neutrinos from the trajectories of their parents[28].

Disappearance only.- Consider the case of no detectable decay products, that is, the neutrinos simply disappear. This limit is interesting for decay to 'invisible' daughters, such as a sterile neutrino, and also for decay to active daughters if the source spectrum falls sufficiently steeply with energy. In the latter case, the flux of daughters of degraded energy will make a negligible contribution to the total flux at a given energy. Since coherence will be lost we have.

$$\phi_{\nu_\alpha} = \sum_{i\beta} \phi_{\nu_\beta}^{\text{source}}(E) |U_{\beta i}|^2 |U_{\alpha i}|^2 e^{-L/\tau_i(E)} \quad (11)$$

$$\xrightarrow{L \gg \tau_i} \sum_{i(\text{stable}), \beta} \phi_{\nu_\beta}^{\text{source}}(E) |U_{\beta i}|^2 |U_{\alpha i}|^2, \quad (12)$$

where the ϕ_{ν_α} are the fluxes of ν_α , $U_{\alpha i}$ are elements of the neutrino mixing matrix and τ are

the neutrino lifetimes in the laboratory frame. Eq. (5) corresponds to the case where decay is complete by the time the neutrinos reach us, so only the stable states are included in the sum.

The simplest case (and the most generic expectation) is a normal hierarchy in which both ν_3 and ν_2 decay, leaving only the lightest stable eigenstate ν_1 . In this case the flavor ratio is $U_{e1}^2 : U_{\mu 1}^2 : U_{\tau 1}^2$ [25]. Thus if $U_{e3} = 0$

$$\phi_{\nu e} : \phi_{\nu \mu} : \phi_{\nu \tau} \simeq 5 : 1 : 1, \quad (13)$$

where we used the neutrino mixing parameters given above[8]. Note that this is an extreme deviation of the flavor ratio from that in the absence of decays. It is difficult to imagine other mechanisms that would lead to such a high ratio of ν_e to ν_μ . In the case of inverted hierarchy, ν_3 is the lightest and hence stable state, and so[8]

$$\phi_{\nu e} : \phi_{\nu \mu} : \phi_{\nu \tau} = U_{e3}^2 : U_{\mu 3}^2 : U_{\tau 3}^2 = 0 : 1 : 1. \quad (14)$$

If $U_{e3} = 0$ and $\theta_{atm} = 45^\circ$, each mass eigenstate has equal ν_μ and ν_τ components. Therefore, decay cannot break the equality between the $\phi_{\nu \mu}$ and $\phi_{\nu \tau}$ fluxes and thus the $\phi_{\nu e} : \phi_{\nu \mu}$ ratio contains all the useful information. The effect of a non-zero U_{e3} on the no-decay case of $1 : 1 : 1$ is negligible.

When U_{e3} is not zero, and the hierarchy is normal, it is possible to obtain information on the values of U_{e3} as well as the CPV phase δ [29]. The flavor ratio e/μ varies from 5 to 15 (as U_{e3} goes from 0 to 0.2) for $\cos \delta = +1$ but from 5 to 3 for $\cos \delta = -1$. The ratio τ/μ varies from 1 to 5 ($\cos \delta = +1$) or 1 to 0.2 ($\cos \delta = -1$) for the same range of U_{e3} .

If the decays are not complete and if the daughter does not carry the full energy of the parent neutrino; the resulting flavor mix is somewhat different but any case it is still quite distinct from

the simple $1 : 1 : 1$ mix[8].

Incidentally, neutrino decay also affects the signals for relic supernova $\bar{\nu}_e$ s and the sensitivity extends to 10^{10} sec/eV. The main results can be summarized as follows[30,31]. If we assume complete decay as before (for simplicity), then for normal hierarchy, the signal is enhanced by about a factor of 2; and for inverted hierarchy, the signal goes away.

6. Magnetic Moments and Other Neutrino Properties

If the path of neutrinos takes them thru regions with significant magnetic fields and the neutrino magnetic moments are large enough, the flavor mix can be affected[32]. The main effect of the passage thru magnetic field is the conversion of a given helicity into an equal mixture of both helicity states. This is also true in passage thru random magnetic fields[33].

If the neutrino are Dirac particles, and all magnetic moments are comparable, then the effect of the spin-flip is to simply reduce the overall flux of all flavors by half, the other half becoming the sterile Dirac partners.

If the neutrinos are Majorana particles, the flavor composition remains $1 : 1 : 1$ when it starts from $1 : 1 : 1$, and the absolute flux remains unchanged.

What happens when large magnetic fields are present in or near the neutrino production region? In case of Dirac neutrinos, there is no difference and the outcoming flavor ratio remains $1 : 1 : 1$, with the absolute fluxes reduced by half. In case of Majorana neutrinos, since the initial flavor mix is no longer universal but is $\nu_e : \nu_\mu : \nu_\tau \approx 1 : 2 : 0$, this is modified but it turns out that the final(post-oscillation) flavor mix is still $1 : 1 : 1$!

As for mixing with sterile neutrinos, if the mixings are small, there are small deviations from the universality[7]. A specific case of large mixing and very small δm^2 is discussed in the next section.

Other neutrino properties can also affect the neutrino flavor mix and modify it from the canonical $1 : 1 : 1$. If neutrinos have flavor(and equivalence principle) violating couplings to gravity(FVG), or Lorentz invariance violating(CPT violating or conserving) couplings; then there can be resonance effects which make for one way transitions(analogues of MSW transitions) e.g. $\nu_\mu \rightarrow \nu_\tau$ but not vice versa[34,35]. In case of FVG for example, this can give rise to an anisotropic deviation of the ν_μ/ν_τ ratio from 1, becoming less than 1 for events coming from the direction towards the Great Attractor, while remaining 1 in other directions[34].

Another possibility that can give rise to deviations of the flavor mix from the canonical $1 : 1 : 1$ is the idea of neutrinos of varying mass(MaVaNs). In this proposal[36], by having the dark energy and neutrinos(a sterile one to be specific) couple, and track each other; it is possible to relate the small scale 2×10^{-3} eV required for the dark energy to the small neutrino mass, and furthermore the neutrino mass depends inversely on neutrino density, and hence on the epoch. As a result, if this sterile neutrino mixes with a flavor neutrino, the mass difference varies along the path, with potential resonance enhancement of the transition probability into the sterile neutrino, and thus change the flavor mix[37]. For example, if only one resonance is crossed enroute, it can lead to a conversion of the lightest (mostly) flavor state into the (mostly) sterile state, thus changing the flavor mix to $1 - U_{e1}^2 : 1 - U_{\mu 1}^2 : 1 - U_{\tau 1}^2 \approx 1/3 : 1 : 1$, in case of normal hierarchy and similarly

$\approx 2 : 1 : 1$ in case of inverted hierarchy.

7. Pseudo-Dirac Neutrinos with very small mass differences [38]

If each of the three neutrino mass eigenstates is actually a doublet with very small mass difference (smaller than $10^{-6}eV$), then there are no current experiments that could have detected this. Such a possibility was raised long ago[39]. It turns out that the only way to detect such small mass differences ($10^{-12}eV^2 > \delta m^2 > 10^{-18}eV^2$) is by measuring flavor mixes of the high energy neutrinos from cosmic sources. Fig. 1 shows that relic supernova neutrino signals and AGN neutrinos are sensitive to mass difference squared down to $10^{-20}eV^2$.

Let $(\nu_1^+, \nu_2^+, \nu_3^+; \nu_1^-, \nu_2^-, \nu_3^-)$ denote the six mass eigenstates where ν^+ and ν^- are a nearly degenerate pair. A 6x6 mixing matrix rotates the mass basis into the flavor basis $(\nu_e, \nu_\mu, \nu_\tau; \nu_e, \nu_\mu, \nu_\tau)$. In general, for six Majorana neutrinos, there would be fifteen rotation angles and fifteen phases. However, for pseudo-Dirac neutrinos, Kobayashi and Lim[40] have given an elegant proof that the 6x6 matrix V_{KL} takes the very simple form (to lowest order in $\delta m^2/m^2$):

$$V_{KL} = \begin{pmatrix} U & 0 \\ 0 & U_R \end{pmatrix} \cdot \begin{pmatrix} V_1 & iV_1 \\ V_2 & -iV_2 \end{pmatrix}, \quad (15)$$

where the 3×3 matrix U is just the usual mixing matrix determined by the atmospheric and solar observations, the 3×3 matrix U_R is an unknown unitary matrix and V_1 and V_2 are the diagonal matrices $V_1 = \text{diag}(1, 1, 1)/\sqrt{2}$, and $V_2 = \text{diag}(e^{-i\phi_1}, e^{-i\phi_2}, e^{-i\phi_3})/\sqrt{2}$, with the ϕ_i being arbitrary phases.

As a result, the three active neutrino states are described in terms of the six mass eigenstates as:

$$\nu_{\alpha L} = U_{\alpha j} \frac{1}{\sqrt{2}} (\nu_j^+ + i\nu_j^-). \quad (16)$$

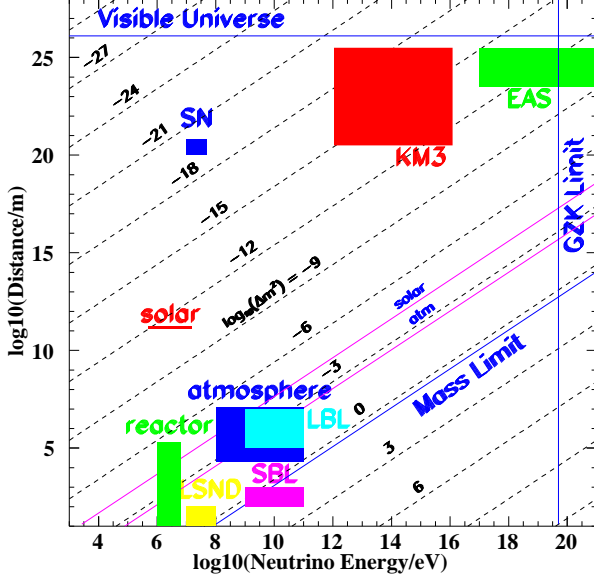


Figure 1. The ranges of distance and energy covered in various neutrino experiments. The diagonal lines indicate the mass-squared differences (in eV^2) that can be probed with vacuum oscillations; at a given L/E , larger δm^2 values can be probed by averaged-out oscillations. We focus on a neutrino telescope of 1-km scale (denoted “KM3”), or larger, if necessary. From Ref.[28].

The nontrivial matrices U_R and V_2 are not accessible to active flavor measurements. The flavor conversion probability can thus be expressed as

$$P_{\alpha\beta} = \frac{1}{4} \left| \sum_{j=1}^3 U_{\alpha j} \left\{ e^{i(m_j^+)^2 L/2E} + e^{i(m_j^-)^2 L/2E} \right\} U_{\beta j}^* \right|^2$$

The flavor-conserving probability is also given by this formula, with $\beta = \alpha$. Hence, in the description of the three active neutrinos, the only new parameters are the three pseudo-Dirac mass differences, $\delta m_j^2 = (m_j^+)^2 - (m_j^-)^2$. In the limit that they are negligible, the oscillation formulas

Table 1

Flavor ratios $\nu_e : \nu_\mu$ for various scenarios. The numbers j under the arrows denote the pseudo-Dirac splittings, δm_j^2 , which become accessible as L/E increases. Oscillation averaging is assumed after each transition j . We have used $\theta_{\text{atm}} = 45^\circ$, $\theta_{\text{solar}} = 30^\circ$, and $U_{e3} = 0$.

1 : 1	$\xrightarrow{3}$	4/3:1	$\xrightarrow{2,3}$	14/9 : 1	$\xrightarrow{1,2,3}$	1 : 1
1 : 1	$\xrightarrow{1}$	2/3:1	$\xrightarrow{1,2}$	2/3 : 1	$\xrightarrow{1,2,3}$	1 : 1
1 : 1	$\xrightarrow{2}$	14/13:1	$\xrightarrow{2,3}$	14/9 : 1	$\xrightarrow{1,2,3}$	1 : 1
1 : 1	$\xrightarrow{1}$	2/3:1	$\xrightarrow{1,3}$	10/11 : 1	$\xrightarrow{1,2,3}$	1 : 1
1 : 1	$\xrightarrow{3}$	4/3:1	$\xrightarrow{1,3}$	10/11 : 1	$\xrightarrow{1,2,3}$	1 : 1
1 : 1	$\xrightarrow{2}$	14/13:1	$\xrightarrow{1,2}$	2/3 : 1	$\xrightarrow{1,2,3}$	1 : 1

reduce to the standard ones and there is no way to discern the pseudo-Dirac nature of the neutrinos.

L/E -Dependent Flavor Ratios.— Given the enormous pathlength between astrophysical neutrino sources and the Earth, the phases due to the relatively large solar and atmospheric mass-squared differences will average out (or equivalently, decohere). The probability for a neutrino telescope to measure the flavor ν_β is then:

$$P_\beta = \sum_\alpha w_\alpha \sum_{j=1}^3 |U_{\alpha j}|^2 |U_{\beta j}|^2 \left[1 - \sin^2 \left(\frac{\delta m_j^2 L}{4E} \right) \right]$$

where w_α represents the fraction of the flavor α present initially. In the limit that $\delta m_j^2 \rightarrow 0$, the expression above reproduces the standard form. The new oscillation terms are negligible until E/L becomes as small as the tiny pseudo-Dirac mass-squared splittings δm_j^2 .

The flavors deviate from the democratic value of $\frac{1}{3}$ by

$$\begin{aligned} \delta P_e &= -\frac{1}{3} \left[\frac{3}{4} \chi_1 + \frac{1}{4} \chi_2 \right], \\ \delta P_\mu = \delta P_\tau &= -\frac{1}{3} \left[\frac{1}{8} \chi_1 + \frac{3}{8} \chi_2 + \frac{1}{2} \chi_3 \right] \end{aligned}$$

where $\chi_i = \sin^2(\delta m_i^2 L/4E)$.

Table 1 shows how the $\nu_e : \nu_\mu$ ratio is altered if we cross the threshold for one, two, or all three of the pseudo-Dirac oscillations. The flavor ratios deviate from 1 : 1 when one or two of the pseudo-Dirac oscillation modes is accessible. In the ultimate limit where L/E is so large that all three oscillating factors have averaged to $\frac{1}{2}$, the flavor ratios return to 1 : 1, with only a net suppression of the measurable flux, by a factor of $1/2$.

8. Cosmology with Neutrinos

If the oscillation phases can indeed be measured for the very small mass differences by the deviations of the flavor mix from 1 : 1 : 1 as discussed above, the following possibility is raised. It is a fascinating fact that non-averaged oscillation phases, $\delta\phi_j = \delta m_j^2 t/4p$, and hence the factors χ_j , are rich in cosmological information[26,41]. Integrating the phase backwards in propagation time, with the momentum blue-shifted, one obtains

$$\delta\phi_j = \int_0^{z_e} dz \frac{dt}{dz} \frac{\delta m_j^2}{4p_0(1+z)} \quad (17)$$

$$= \left(\frac{\delta m_j^2 H_0^{-1}}{4p_0} \right) I \quad (18)$$

where I is given by

$$I = \int_1^{1+z_e} \frac{d\omega}{\omega^2} \frac{1}{\sqrt{\omega^3 \Omega_m + (1 - \Omega_m)}}, \quad (19)$$

z_e is the red-shift of the emitting source, and H_0^{-1} is the Hubble time, known to 10% [42]. This result holds for a flat universe, where $\Omega_m + \Omega_\Lambda = 1$, with Ω_m and Ω_Λ the matter and vacuum energy densities in units of the critical density. The integral I is the fraction of the Hubble time available for neutrino transit. For the presently preferred values $\Omega_m = 0.3$ and $\Omega_\Lambda = 0.7$, the asymptotic

($z_e \rightarrow \infty$) value of the integral is 0.53. This limit is approached rapidly: at $z_e = 1$ (2) the integral is 77% (91%) saturated. For cosmologically distant ($z_e > 1$) sources such as gamma-ray bursts, non-averaged oscillation data would, in principle, allow one to deduce δm^2 to about 20%, without even knowing the source red-shifts. Known values of Ω_m and Ω_Λ might allow one to infer the source redshifts z_e , or vice-versa.

This would be the first measurement of a cosmological parameter with particles other than photons. An advantage of measuring cosmological parameters with neutrinos is the fact that flavor mixing is a microscopic phenomena and hence presumably free of ambiguities such as source evolution or standard candle assumptions[26,43]. Another method of measuring cosmological parameters with neutrinos is given in Ref.[44].

9. Experimental Flavor Identification

It is obvious from the above discussion that flavor identification is crucial for the purpose at hand. In a water cerenkov detector flavors can be identified as follows.

The ν_μ flux can be measured by the μ 's produced by the charged current interactions and the resulting μ tracks in the detector which are long at these energies. ν_e 's produce showers by both CC and NC interactions. The total rate for showers includes those produced by NC interactions of ν_μ 's and ν_τ 's as well and those have to be (and can be) subtracted off to get the real flux of ν_e 's. However, this distinction between hadronic showers of neutral current events and the electron-containing charged current events is rather difficult to make. Double-bang and lollipop events are signatures unique to tau neutrinos, made possible by the fact that tau leptons decay before they lose a

significant fraction of their energy. Double-bang events consists of a hadronic shower initiated by a charged-current interaction of the ν_τ followed by a second energetic shower (hadronic or electromagnetic) from the decay of the resulting tau lepton[4]. Lollipop events consist of the second of the double-bang showers along with the reconstructed tau lepton track (the first bang may be detected or not). In principle, with a sufficient number of events, a fairly good estimate of the flavor ratio $\nu_e : \nu_\mu : \nu_\tau$ can be reconstructed, as has been discussed recently. Deviations of the flavor ratios from $1 : 1 : 1$ due to possible decays are so extreme that they should be readily identifiable[45]. Upcoming high energy neutrino telescopes, such as Icecube[46], will not have perfect ability to separately measure the neutrino flux in each flavor. However, the quantities we need are closely related to observables, in particular in the limit of $\nu_\mu - \nu_\tau$ symmetry ($\theta_{atm} = 45^\circ$ and $U_{e3} = 0$), in which all mass eigenstates contain equal fractions of ν_μ and ν_τ . In that limit, the fluxes for ν_μ and ν_τ are always in the ratio $1 : 1$, with or without decay. This is useful since the ν_τ flux is the hardest to measure.

Even in the extreme case when one assumes that tau events are not identifiable, something about the flavor mix can be deduced. Let the only experimental information available be the number of muon tracks and the number of showers. The relative number of shower events to track events can be related to the most interesting quantity for testing decay scenarios, i.e., the ν_e to ν_μ ratio. The precision of the upcoming experiments should be good enough to test the extreme flavor ratios produced by decays. If electromagnetic and hadronic showers can be separated, then the precision will be even better[45].

Comparing, for example, the standard flavor ratios of $1 : 1 : 1$ to the possible $5 : 1 : 1$ generated by decay, the more numerous electron neutrino flux will result in a substantial increase in the relative number of shower events.

The details of this observation depends on the range of muons generated in or around the detector and the ratio of charged to neutral current cross sections. The measurement will be limited by the energy resolution of the detector and the ability to reduce the atmospheric neutrino background. The atmospheric background drops rapidly with energy and should be negligibly small at and above the PeV scale.

10. Discussion and Conclusions

The flux ratios we discuss are energy-independent because we have assumed that the ratios at production are energy-independent, that all oscillations are averaged out, and that all possible decays are complete. In the standard scenario with only oscillations, the final flux ratios are $\phi_{\nu_e} : \phi_{\nu_\mu} : \phi_{\nu_\tau} = 1 : 1 : 1$. In the cases with decay, we have found rather different possible flux ratios, for example $5 : 1 : 1$ in the normal hierarchy and $0 : 1 : 1$ in the inverted hierarchy. These deviations from $1 : 1 : 1$ are so extreme that they should be readily measurable.

If we are very fortunate[47], we may be able to observe a reasonable number of events from several sources (of known distance) and/or over a sufficient range in energy. Then the resulting dependence of the flux ratio (ν_e/ν_μ) on L/E as it evolves from say 5 (or 0) to 1, can be clear evidence of decay and further can pin down the actual lifetime instead of just placing a bound.

To summarize, we suggest that if future measurements of the flavor mix at earth of high en-

ergy astrophysical neutrinos find it to be

$$\phi_{\nu_e}/\phi_{\nu_\mu}/\phi_{\nu_\tau} = \alpha/1/1; \quad (20)$$

then

- (i) $\alpha \approx 1$ (the most boring case) confirms our knowledge of the MNSP[2] matrix and our prejudice about the production mechanism;
- (ii) $\alpha \approx 1/2$ indicates that the source emits pure ν'_μ s and the mixing is conventional;
- (iii) $\alpha \approx 3$ from a unique direction, e.g. the Cygnus region, would be evidence in favour of a pure $\bar{\nu}_e$ production as has been suggested recently[48];
- (iv) $\alpha > 1$ indicates that neutrinos are decaying with normal hierarchy; and
- (v) $\alpha \ll 1$ would mean that neutrino decays are occuring with inverted hierarchy;
- (vi) Values of α which cover a broader range (3 to 15) and deviation of the μ/τ ratio from 1 (between 0.2 to 5) can yield valuable information about U_{e3} and $\cos \delta$. Deviations of α which are less extreme (between 0.7 and 1.5) can also probe very small pseudo-Dirac δm^2 (smaller than $10^{-12} eV^2$).

Incidentally, in the last three cases, the results have absolutely no dependence on the initial flavor mix, and so are completely free of any dependence on the production model. So either one learns about the production mechanism and the initial flavor mix, as in the first three cases, or one learns only about the neutrino properties, as in the last three cases. In any case, it should be evident that the construction of very large neutrino detectors is a “no lose” proposition.

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